

**Summary of Research**  
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**MODELLING ATMOSPHERIC SCATTERERS USING SPACECRAFT  
OBSERVATIONS**

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Submitted to  
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## Modelling Atmospheric Scatterers Using Spacecraft Observations

The period covered by this cooperative agreement included the analysis of data from the Voyager encounter with Neptune and Triton and the primary Galileo mission to Jupiter (including the Galileo Probe entry into Jupiter's atmosphere), as well as continued work on Uranus' seasonal variability using the Voyager encounter data as a baseline.

### Analysis of Voyager Triton Limb Images ...

High resolution images of planetary limbs can be used to derive information about the vertical structure and single scattering properties of scatterers in the upper atmosphere (assuming, of course, that the planet has an atmosphere). Because of the very long slant path through the atmosphere of lines of sight grazing the edge of the planet, such observations generally reveal information about altitudes corresponding to optical depths in the range  $10^{-5}$  to  $\sim 0.03$ . In this respect limb studies complement observations of the disk of the planet, which are most sensitive to structure and composition at optical depths of 0.1 to  $\sim 3$ . Also, extinction as a function of altitude can be derived in a more-or-less direct fashion from a limb intensity scan. If there are any discrete scattering layers present they can be observed directly instead of being inferred from (probably ambiguous) fits of atmospheric scattering models to observations of specific intensity (I/F) from the planet as a whole (Pollack *et al.* 1987; Rages and Pollack 1983).

Voyager images of Triton's limb reveal the presence of scattering clouds and hazes above the surface, with considerable latitudinal variability (Pollack *et al.* 1990). Standard algorithms for calculating scattered specific intensities (I/F's) near the limb of a spherical atmosphere were combined with Mie theory to derive models for the single scattering properties and the vertical distribution of these clouds and hazes. Near  $14^{\circ}\text{S}$  latitude,  $275^{\circ}\text{W}$  longitude only haze was present, with optical depth  $\sim 0.01$  in violet and  $\sim 0.004$  in orange. Near  $52^{\circ}\text{S}$  latitude,  $124^{\circ}\text{W}$  longitude condensation clouds were also present below 6 km, with optical depth  $\geq 0.1$  and mean particle radii either near  $0.25\text{ }\mu\text{m}$  or near  $0.7\text{ }\mu\text{m}$  (Rages and Pollack 1992).

### ... and Neptune Limb Images ...

Voyager 2 obtained a high-phase-angle clear-filter image of Neptune's limb at  $\sim 30^{\circ}\text{S}$  latitude with 2 km spatial resolution. Inversion of the radial intensity profile across the limb reveals a rapid increase in extinction with decreasing altitude near the 12-mbar pressure level, about where photochemically produced ethane would be expected to condense. At about 30 mbar, the haze particle radius and number density are found to be  $0.13 \pm 0.02\text{ }\mu\text{m}$  and  $5 \pm 3\text{ cm}^{-3}$ , respectively. The total haze extinction optical depth is  $\sim 0.003$  above 15 mbar, and is mostly due to scattering rather than absorption, which makes heating due to absorption of sunlight by the haze inadequate to account for Neptune's observed stratospheric temperature inversion (Baines *et al.* 1993; Moses *et al.* 1995).

The most prominent albedo feature on Neptune at the time of the Voyager encounter was the Great Dark Spot and accompanying Bright Companion at  $\sim 27^{\circ}\text{S}$ . There has been speculation that the cores of features such as the Bright Companion might be convective columns with tops punching through the "cold trap" of the tropopause at  $\sim 150$  mbar, and as such might serve as sources for the observed excess of methane gas in the stratosphere. A comparison of the center-to-limb brightness curves of the Bright Companion in the Voyager wide-angle filters with those of its unperturbed surroundings revealed that, despite its high contrast in the methane absorption band at 619 nm, the top of the Bright Companion may lie as deep as the 400 mbar pressure level. This is well below the tropopause, and would eliminate the Bright Companion and similar features as sources for the observed excess methane mixing ratio in the stratosphere (Rages *et al.* 1996).

### **... and Galileo Images of Venus' Limb ...**

The Galileo spacecraft flew past Venus on its roundabout trajectory to Jupiter in February 1990. Imaging data included six images of the limb between  $\sim 10^\circ\text{N}$  and  $\sim 60^\circ\text{N}$  (three pairs consisting of one violet image and one  $1\text{-}\mu\text{m}$  image each). These images were inverted to yield vertical extinction profiles using the previously developed procedure. A region of enhanced extinction was clearly visible in both filters at the highest latitude and an altitude of 91 km (pressure  $\approx 0.3$  mbar). At  $48^\circ\text{N}$  a similar feature was visible in the  $1\text{-}\mu\text{m}$  filter at an altitude of 88 km ( $\approx 0.6$  mbar) and there was also some enhancement in the violet extinction at this altitude, although it was far less conspicuous. Near the equator there was a change in the mean scale height near 87 km (1 bar), but no sign of any other feature. The fact that the extinction at a given altitude is much the same at  $0.41\text{-}\mu\text{m}$  (violet) and  $1\text{-}\mu\text{m}$  indicates that the dominant sources of scattered light at all three latitudes are probably greater than  $0.1\text{-}0.2\text{-}\mu\text{m}$  in size. Particles smaller than this would be Rayleigh scattering, and the extinction coefficient for Rayleigh scatterers would be expected to drop by more than an order of magnitude between  $0.41\text{-}\mu\text{m}$  and  $1\text{-}\mu\text{m}$ . A comparison of the extinction profiles at  $57^\circ\text{N}$  and  $48^\circ\text{N}$  gives the impression that the layer of enhanced extinction may drop to lower altitudes at lower latitudes. (Belton *et al.* 1991).

### **... and (finally) Jupiter's Limb**

In December 1996, during orbit E4, Galileo used its Solid State Imaging subsystem (SSI) to take four images of Jupiter's equatorial limb and four images of the limb in the transition zone between north temperate latitudes and the northern polar region. Galileo made images of each region in violet (mean wavelength 411 nm) and near infrared (756 nm) filters at two high solar phase angles. The images were taken close enough to Jupiter to give spatial resolutions of  $\sim 15$  km/pixel, comparable to a scale height near the tropopause. They reveal the presence of stratospheric hazes in the equatorial region ( $9^\circ\text{N}$ ) and the north polar transition zone ( $60^\circ\text{N}$ ). A discrete layer detached from the limb is present at  $60^\circ\text{N}$ ,  $315^\circ\text{W}$ , but not  $20^\circ$  further east at  $60^\circ\text{N}$ ,  $285^\circ\text{W}$ . Bright streaks running roughly north-south are also present on Jupiter's crescent at  $60^\circ\text{N}$ . No such discrete features are seen at  $9^\circ\text{N}$ , where the haze appears to be more uniformly distributed in height and longitude.

Radial specific intensity (I/F) profiles across the limb were inverted to give vertical extinction profiles over  $\sim 200$  km in Jupiter's stratosphere at both  $9^\circ\text{N}$  and  $60^\circ\text{N}$ . The maximum I/F in each image was used to estimate the single scattering phase function for the corresponding scattering angle, wavelength, and altitude. These phase functions required different particle sizes in the violet and near-IR filters, which probed two different altitudes. Haze distribution models were found for both latitudes in which the near-IR images constrain the haze properties near or above the 100 mbar pressure level, where the mean particle radius is about  $0.45\text{-}\mu\text{m}$  and the haze number density is near  $0.15\text{ cm}^{-3}$ . In these models the violet images constrain the haze properties near or above the 20 mbar pressure level, where the haze particle size is about  $0.32\text{-}\mu\text{m}$  at  $9^\circ\text{N}$  and about  $0.27\text{-}\mu\text{m}$  at  $60^\circ\text{N}$ . At this altitude, the haze number density increases by almost an order of magnitude between the equator and the polar transition region. To fit Hubble Space Telescope (HST) observations in the ultraviolet and the methane absorption band at 890 nm, these models must be modified by the addition of another scattering component, consisting of  $\leq 0.02\text{-}\mu\text{m}$  particles which are conservatively scattering in the near-IR and have single scattering albedos near 0.5 in the UV. This  $0.02\text{-}\mu\text{m}$  haze accounts for more than 90% of the Rayleigh scattering (originally ascribed to scattering by the hydrogen-helium gas) in both filters, with a corresponding drop of a factor of  $\sim 10$  in the pressure levels probed (Rages *et al.* 1999).

### **Latitude Bands and Temporal Variability on Uranus**

Uranus undergoes extreme seasonal variations as a result of its  $98^\circ$  axial tilt. Since the Voyager encounter in January 1986, the subsolar latitude has shifted from its southern solstice position of  $82^\circ\text{S}$  to about  $30^\circ\text{S}$ . During the past 20 years, Uranus' brightness has varied by about 4% in blue

light (wavelength  $0.472 \pm 0.021 \mu\text{m}$ ) and about 11% in yellow ( $0.551 \pm 0.023 \mu\text{m}$ ) (Lockwood *et al.* 1985; Lockwood, private communication). These changes appear to be seasonal rather than being related to solar variability, which would show up as an  $\sim 11$  year periodicity. Some part of this seasonal variability may be due simply to the changing aspect of the planet as seen from Earth, since Voyager found that southern polar latitudes were somewhat brighter than regions nearer the equator.

Radiative transfer models were derived for five southern latitudes in addition to the two ( $22.5^\circ\text{S}$  and  $65^\circ\text{S}$ ) previously studied using Voyager images (Rages *et al.* 1991). Both  $\tau$ , the optical depth of Uranus' methane cloud layer, and  $f_H$ , the product of the mean methane cloud particle size and the fraction of stratospheric haze material mixed into the methane cloud, vary by factors of  $\sim 2$  between the latitudes of  $10^\circ\text{S}$  and  $85^\circ\text{S}$ . The variability observed by Lockwood exceeds that which would be caused by these latitudinal variations in atmospheric structure, by about a factor of two in the blue and about a factor of five in the yellow. Therefore Uranus is undergoing actual temporal changes in its scattering properties. Calculations show that even the most favorable combination of changes in the methane cloud optical depth and the concentration of stratospheric haze and its residue below the methane cloud produce only about  $2/3$  of the observed brightness changes in both filters. Some other factor, such as hemispheric asymmetry in the atmospheric structure, must be introduced to account for Uranus' observed long-term albedo changes (Rages and Pollack 1994).

### Galileo Probe Nephelometer

The Galileo Probe entered Jupiter's atmosphere on December 7, 1995 and returned data for 57 minutes, down to a pressure level of about 20 bars. One of the instruments carried on the probe was a nephelometer intended to measure the scattering properties of the regions the Probe passed through. The Probe entered Jupiter's atmosphere near the edge of a  $5\text{-}\mu\text{m}$  hot spot, a region where the atmosphere is much clearer than on most of the planet. The nephelometer definitely saw one thin cloud at about 1.3 bars, and very preliminary analysis indicates that this cloud is composed of particles with mean radii of  $1\text{-}5 \mu\text{m}$ . The data can be fit either by spherical scatterers having radii of  $0.8 \mu\text{m}$ , a real refractive index  $n_r$  of  $1.7\text{-}1.8$ , and an imaginary refractive index  $n_i \sim 0.05$ , or by scatterers with radii around  $8 \mu\text{m}$ , a very low value of  $n_r$  ( $\leq 1.2$ ), and  $n_i = 0$ . The high- $n_r$  solution is consistent with the little that is known of the optical constants of ammonium hydrosulfide, the most likely constituent of a cloud at 1.3 bars, but would require the presence of some additional component to account for the absorption. The low- $n_r$  solution requires either a very unusual composition or "fluffy" particles with many internal voids, since none of the likely bulk constituents have refractive indices below about 1.3. The cloud optical depth is near 2, over a vertical distance of  $\sim 20$  km (Ragent *et al.* 1996; Ragent *et al.* 1998).

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## **APPENDIX**

### **Subject Inventions Certification**

There were no subject inventions required to be disclosed to NASA which resulted from this work. There were no subcontracts awarded under this Cooperative Agreement.

